

If R large, what is M_p ?

Economy: $M_p \sim \text{TeV}$

Then
$$R \sim \frac{1}{M_p} \cdot \left(\frac{M_4}{M_p} \right)^{\frac{2}{n}}$$

$$\begin{array}{ccc} \uparrow & & \uparrow \\ 10^{-17} \text{ cm} & & 10^{16} \end{array}$$

$$n = 1 \quad \times$$

$$n = 2 \quad R \sim \text{mm}$$

⋮

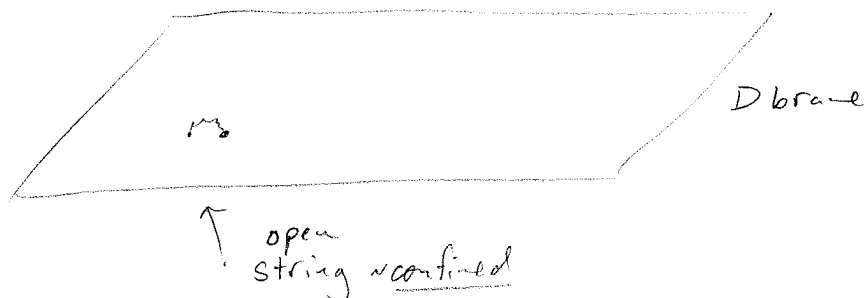
$$n = 6 \quad R \sim 10^{-12} \text{ cm} \quad (\text{Strings?})$$

Certifiably crazy: nucleus $\sim 10^{-13} \text{ cm}$
 $M_W^{-1} \sim 10^{-16} \text{ cm}$

extra dims not seen.
 (eg. mods to $\frac{1}{r}$ pots, ...)


But: \exists gravitational "defects" \sim domain walls,

in particular ST \rightsquigarrow D-branes.

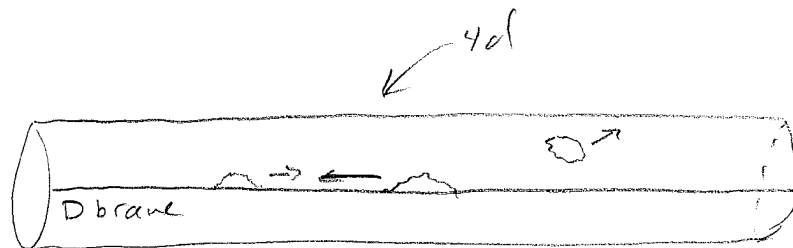


So: if SM matter = open strings,
 gauge

they behave as if in fewer dims

gravity:  closed strings; higher dimensional

Picture



Detailed Q's :

Realizable geometry (moduli issues)
 SM gauge content
 $SU(2) \times U(1)$ breaking
 GUT relⁿ
 \rightarrow masses
 no p decay
 ...

looks challenging, some real progress.

Actually warped compactifications look better

\leftarrow scale varies over y

$$ds^2 = e^{2A(y)} dx_4^2 + ds_6^2(y)$$

$$S_{\text{grav}} \sim M_P^8 \int d^4x \sqrt{g_4} \int d^6y \sqrt{g_6} e^{+2A} \mathcal{R}_4 + \dots$$

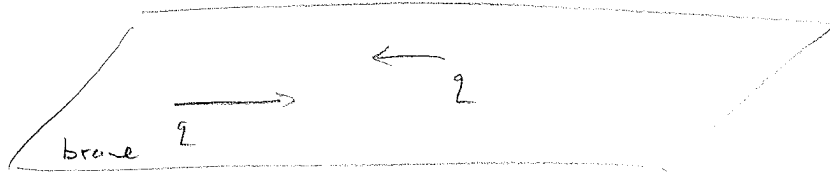
(ex-scaling)

$$\therefore \frac{M_4^2}{M_P^2} \sim M_P^6 \underbrace{\int d^6y \sqrt{g_6} e^{2A}}_{\text{warped volume}}$$

large warping \Leftrightarrow large M_4 .

Can we rule this out directly?

Locally:



gauge interactions - $4d$

gravity corrections - higher d

but we're not so sensitive!

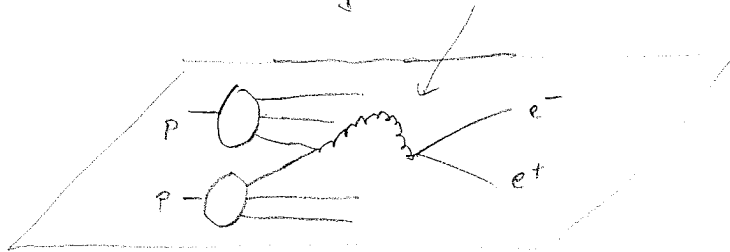
eg. mods to gravity: if $n=2$, $r \sim 1 \mu\text{m}$,
pushed hard by exp.

but $n=6$ not.

accelerators?

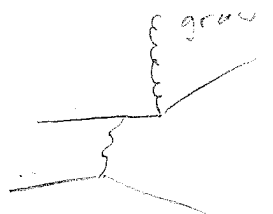
eg Drell Yan

graviton - KK excitation.



$$\sigma \sim \frac{s^2}{M_p^4} \quad ? \quad (\text{spectrum, etc.})$$

or



missing energy.

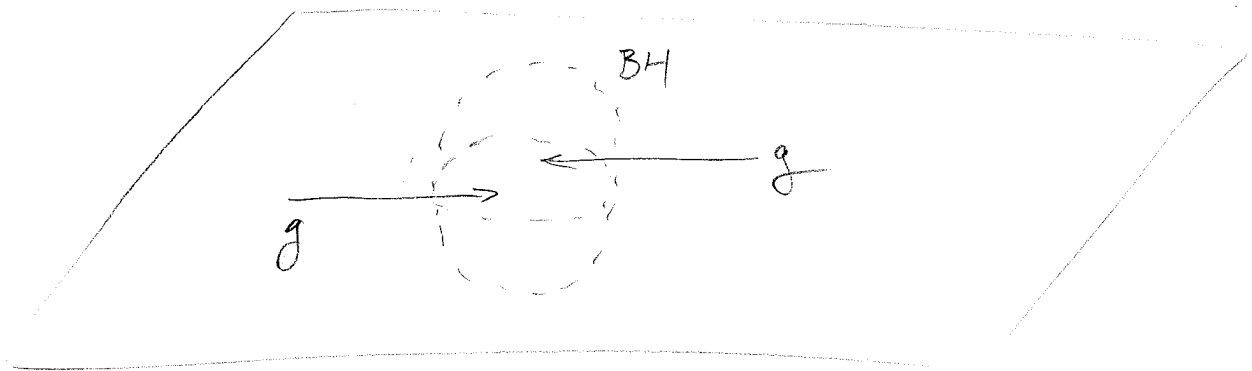
Wia @ high energies

Tevatron bds $\leadsto M_p \approx \text{TeV.} \quad !$

If gravity gets strong, pretty soon

1. Make strings (if S.T. correct) at $E \approx M_s$
2. Make black holes at $E \approx M_p \quad !!!$

Basic picture



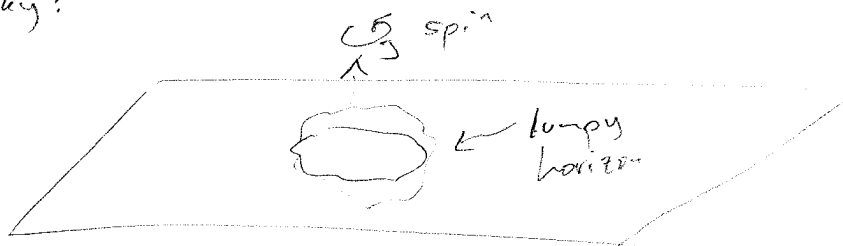
$$\sigma \sim r_{\text{Schw}}^2 \sim \frac{1}{M_p^2} \left(\frac{E}{M_p} \right)^{\frac{2}{1+n}}$$

ie. $\mathcal{O}(\text{TeV}^{-2})$ at threshold!
 (really only black hole if $E \gg M_p$)
 true for partons; energy distributed.

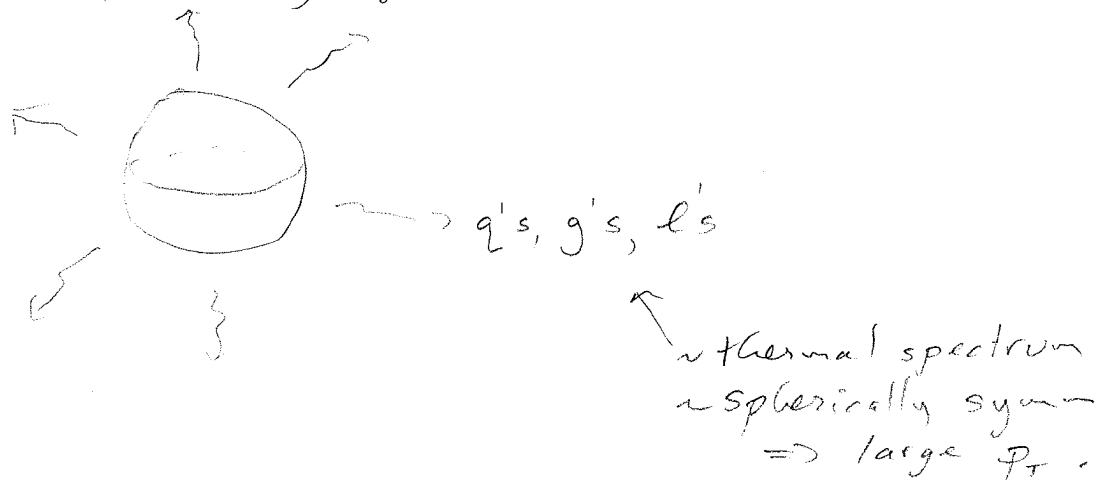
but $M_p \sim 1 \text{ TeV}$, $M_{\text{BH, min}} \sim 5 \text{ TeV}$

$\leadsto \frac{1}{8}$, for example.

Decay:



1. Balding & discharge $\rightarrow M, J$
2. Hawking radiation
 - A. Eliminate spin
 - B. Spherically symmetric - Schwarzschild



End of short distance physics:

$$E \uparrow : \quad \text{QFT} \quad \Delta x \sim \frac{1}{E} \quad \downarrow$$

$$\text{Gravity} \quad R_s \sim E^{\frac{1}{1+n}} \quad \uparrow$$

no longer probe short distances.

An appropriate ending for this course...
we'll see what LHC brings!